



ASTRO DIVER
EXPLORE TWO WORLDS

OSC SHO PROCESSING WORKFLOW

with Perfect Palette Picker (PPP)

Optolong L-eXtreme + L-Synergy with OSC cameras

Two complementary filters. Three emission-line masters. Many palette interpretations. Your vision.

SAME DATA.
MANY PALETTES.
MORE POSSIBILITIES.

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1 ACQUIRE DATA
Capture two complementary dual-band datasets with your OSC camera.



Cooled OSC Astronomy Camera
(e.g. ZWO, QHY, Player One, etc.)

L-eXtreme
(Ha + OIII)



L-Synergy
(SII + OIII)



Optional: UV/IR broadband data for natural RGB stars (≈ 1-2 hours).

2 CALIBRATE & INTEGRATE SEPARATELY
Process each filter's data independently in WBPP.

L-eXtreme Master
(Ha + OIII)

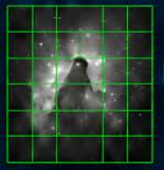


L-Synergy Master
(SII + OIII)



- Use matching calibration files (bias/darks/flats).
- Check FWHM, eccentricity, etc.
- Don't judge L-Synergy by its color – look for the SII signal.

3 REGISTER & ALIGN
Align the two masters pixel-for-pixel.



- Use StarAlignment (if needed).
- Crop to common field.
- Ensure precise alignment for accurate extraction.

4 LINEAR PREPROCESSING
Keep both masters linear.

- ImageSolve (if needed)
- Gradient correction (GraXpert / GX)
- BlurXTerminator (conservative)
- NoiseXTerminator (conservative)
- Preserve weak SII structure

Goal: clean, well-aligned, linear masters ready for DBExtract.

5 EXTRACT EMISSION LINES
Use Seti Astro DBExtract in PixInsight.



- Inputs: registered, linear masters
- OIII is present in both filters and can be used from both datasets
- Generate and save three grayscale masters: Ha, OIII, SII

6 INSPECT EMISSION-LINE MASTERS
Use STF for inspection.





Ha OIII SII

- Ha: typically strongest signal
- OIII: different structures / ionization regions
- SII: often weaker – evaluate S/N and structure

Decision point: add more L-Synergy data if the SII master justifies it.

7 EXPLORE PALETTES WITH PERFECT PALETTE PICKER (PPP)
Use your Ha, OIII and SII masters to explore different palette interpretations.



Compare multiple palettes and choose the one that best reveal the structure and aesthetic you want. Preserve your Ha/OIII/SII masters for future reprocessing.

8 VERIFY PPP OUTPUT STATE (IMPORTANT)
Reset/disable STF and inspect the PPP palette with no screen stretch.

VISIBLE WITHOUT STF



Treat as NONLINEAR

- Already stretched
- Continue with refinement (curves, saturation, contrast, etc.)

BLACK OR NEARLY BLACK



Treat as LINEAR

- Needs a permanent stretch
- Apply HT and/or GHS
- Then continue with nonlinear processing

Do not assume the PPP output is linear or nonlinear based on the setting alone.

CHOOSE A FINISHING APPROACH

A SIMPLIFIED FINISH (Optional)
A fast, straightforward approach. No star separation required.





1. Refine

- Curves / color
- HistogramTransformation (if needed)
- Slight noise reduction

2. Adjust

- Color, contrast, background level
- Minor sharpening (if needed)

3. Final Image

- Check stars and nebula
- Crop
- Ready to share

Best for: a cleaner, simpler workflow or when star separation isn't necessary.

B ADVANCED STARLESS FINISH (Astro Diver Approach)
More control by processing the nebula and stars separately.






1. Separate

- Use StarXTerminator
- Create starless and star images

2. Refine the nebula (according to verified image state)

- CurvesTransformation
- HistogramTransformation
- Masks / selective edits
- Color / contrast adjustments

3. Process stars

- Light adjustments
- Reduce noise (if needed)

4. Recombine

- Combine starless and stars
- Final tweaks
- Background level
- Crop and frame

Best for: maximum control, best aesthetics, and our recommended approach.

9 FINAL RESULT



IC 1396 – The Elephant Trunk Nebula (Foraxx Palette)

Captured with a ZWO ASI533MC Pro cooled OSC camera, Optolong L-eXtreme and L-Synergy filters. Processed in PixInsight using Perfect Palette Picker.

Same data. Many palettes.
The choice is yours.

KEY TAKEAWAYS

- ✓ OSC + complementary dual-band filters give you Ha, OIII and SII data with an OSC camera.
- ✓ Preserve the individual emission-line masters.
- ✓ Perfect Palette Picker lets you explore many interpretations without reacquiring data.
- ✓ Verify PPP output with STF off. If linear, apply a permanent stretch; if nonlinear, continue refining without re-stretching.
- ✓ The workflow is flexible – simple or advanced.
- ✓ The acquisition architecture is more important than committing to one palette in advance.



ASTRO DIVER
EXPLORE TWO WORLDS

CLEARER SKIES.
DEEPER UNDERSTANDING.
BRIGHTER IMAGES.

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ASTRO DIVER | OSC SHO Processing with Perfect Pallet Picker | v1.0

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PG02 — OSC SHO Processing with Perfect Palette Picker

Astro Diver Processing Guide

PURPOSE

This guide documents an Astro Diver workflow for using complementary dual-band one-shot-color data to create and explore synthetic narrowband palettes in PixInsight. The workflow combines an H α + OIII dataset from the Optolong L-eXtreme with an SII + OIII dataset from the Optolong L-Synergy. The datasets remain separate through integration, are aligned precisely, and are then used to create H α , OIII, and SII masters for palette development with Seti Astro tools, including Perfect Palette Picker.

The workflow grew out of processing IC 1396, the Elephant Trunk Nebula. What began as a conventional SHO project evolved into a more flexible approach built around Seti Astro DBExtract and Perfect Palette Picker, allowing several palette interpretations to be evaluated from the same preserved emission-line data.

The goal is to preserve the underlying H α , OIII, and SII information so different synthetic narrowband interpretations can be explored without returning to the original light frames.

OSC TERMINOLOGY

OSC stands for One-Shot Color. An OSC camera is a color astronomy camera that records color information through a Bayer color-filter array in a single exposure, rather than using a monochrome camera with separate color or narrowband filters. The workflow is not tied to a specific camera model. The ZWO ASI533MC Pro is identified later only because it was the OSC camera used for the IC 1396 worked example.

WHY COMPLEMENTARY DUAL-BAND OSC?

A one-shot-color camera records color through a Bayer color-filter array. With an appropriate dual-band filter, that same camera can collect selected emission-line information from an emission nebula. Using two complementary dual-band filters extends the idea: one dataset supplies H α + OIII information and the other supplies SII + OIII information.

For this workflow:

L-eXtreme → H α + OIII

L-Synergy → SII + OIII

Together → H α + OIII + SII information for synthetic narrowband palette development

This provides an OSC route to SHO-style palette exploration while retaining the operational simplicity of a color-camera system. Each dual-band exposure records two selected emission bands, and the extracted line masters can later support several palette interpretations.

OSC AND MONO NARROWBAND — DIFFERENT TRADEOFFS

This approach is not intended to reproduce a dedicated monochrome SHO system exactly. A monochrome camera with separate H α , OIII, and SII filters records each narrowband channel

directly and independently, without a Bayer color-filter array. That architecture generally provides greater channel isolation, signal efficiency, and control over how much integration time is devoted to each emission line.

The complementary dual-band OSC approach trades some of that channel independence for simplicity. It can use an existing cooled color camera, requires only the two complementary filters for this workflow, and records two emission bands during each exposure.

A traditional automated mono narrowband system commonly adds a filter wheel and several high-quality narrowband filters—and often LRGB filters if the rig will also be used for broadband color imaging. The complementary dual-band OSC approach can use an existing cooled color-camera system with fewer acquisition components while still providing a path to synthetic H α /OIII/SII palette processing.

WHAT THE TWO FILTERS CAPTURE

Optolong L-eXtreme — H α + OIII

The L-eXtreme dataset provides the H α and OIII information used by this workflow. In many emission nebulae, H α is the strongest extracted signal and carries substantial nebular structure. OIII can show a different spatial distribution and later contribute cyan, teal, or blue separation depending on the selected palette.

Optolong L-Synergy — SII + OIII

The L-Synergy dataset provides the SII information missing from the L-eXtreme dataset while also recording OIII. SII may be substantially weaker than H α , so the L-Synergy integration should not be judged simply by how impressive its debayered OSC color image appears. The important question is whether extraction reveals useful SII structure at acceptable signal-to-noise.

Why OIII appears twice

OIII is present in both complementary dual-band datasets. A suitable extraction workflow can make use of information from both rather than automatically discarding one OIII dataset. This is one reason the two masters should remain separate and be registered carefully before emission-line extraction.

WHAT H α , OIII, AND SII REPRESENT

H α , OIII, and SII are emission-line signals rather than simply red, green, and blue photographs. H α traces emission from ionized hydrogen and is often especially strong in emission nebulae. OIII traces doubly ionized oxygen and can reveal structures and ionization regions that differ from the hydrogen signal. SII traces singly ionized sulfur and is frequently weaker, but it can reveal additional structures that help distinguish regions within the nebula.

After extraction, each is preserved as a grayscale master. Synthetic narrowband processing assigns those masters to visible colors so their differing spatial distributions can be displayed.

A conventional SHO reference assigns:

SII → Red

H α → Green

OIII → Blue

That is only one possible interpretation. Perfect Palette Picker allows the preserved emission-line information to be explored through alternative palette treatments without reacquiring or restacking the original light frames. In the Elephant project, Foraxx became the preferred final interpretation, but it should be understood as the result of palette evaluation rather than the purpose of the workflow.

WHEN TO USE THIS WORKFLOW

Use this approach when an OSC project contains complementary dual-band observations that collectively provide H α , OIII, and SII information. It is especially useful when the goal is to move beyond a conventional HOO rendering and explore SHO-style or other synthetic narrowband color palettes while retaining the convenience of a one-shot-color imaging system.

The workflow is most useful when all three extracted masters contain meaningful structural information. Weak SII is not automatically failed data, but its signal-to-noise may become the limiting factor and may justify additional L-Synergy integration.

WORKFLOW STARTING POINT

Begin with two calibrated acquisition sets from the same target: one captured through L-eXtreme (H α + OIII) and one through L-Synergy (SII + OIII). The datasets should overlap in framing closely enough to register to a common field. Equipment and acquisition details from the IC 1396 worked example are presented near the end of this guide rather than treated as workflow requirements.

The L-eXtreme and L-Synergy datasets must remain separate through calibration and integration. Do not combine the raw lights from both filters into a single WBPP integration.

1. CALIBRATE AND STACK EACH FILTER SEPARATELY — WBPP

Run the L-eXtreme and L-Synergy lights as separate WBPP stacks with the appropriate calibration frames for each dataset. In PixInsight, the stacking stage is ImageIntegration; here, “stack” is used as the practical term for producing each integrated master. Do not mix lights from the two filters into one stack. The objective is to preserve two independent filter masters that can later be aligned and used for emission-line extraction.

Inspect the input data and integrated masters for tracking failures, clouds, gradients, FWHM/eccentricity problems, and other obvious defects. Do not judge the L-Synergy master only by the appearance of its OSC color image; its important new contribution to this workflow is the SII information that can be extracted from it.

What I'm looking for:

Two clean, independently integrated linear masters with enough usable signal to justify moving to extraction. Weak-looking L-Synergy color is not by itself a reason to reject the dataset.

Decision point:

SII may be substantially weaker than H α . The real test comes after extraction: does the SII master contain useful structure at acceptable signal-to-noise? If not, additional L-Synergy integration may be more valuable than aggressive processing.

2. REGISTER THE TWO MASTERS

The two integrated filter masters must align essentially pixel-for-pixel before emission-line extraction and combination.

Use the stronger or cleaner master as the registration reference. For the Elephant project, the L-eXtreme master is a sensible reference. If necessary, use StarAlignment to register the L-Synergy master to it. If framing or edge artifacts differ between sessions, crop both masters to their common usable field.

What I'm looking for:

Matching geometry across the two masters, with stars and nebular structures coincident and no meaningful framing mismatch that could contaminate later channel extraction or palette construction.

3. CONSERVATIVE LINEAR PREPROCESSING

Keep both masters linear through emission-line extraction and initial palette construction.

ImageSolve when needed so the masters have reliable astrometric information. Correct gradients conservatively with GraXpert/GX when gradients are present, and inspect the corrected result before accepting it. BlurXTerminator can be used conservatively; Correct Only is a sensible starting point before extraction. NoiseXTerminator can also be used when needed, but weak SII deserves particular protection.

The goal is not to make the linear masters look finished. It is to correct problems that would otherwise carry into the extracted channels while preserving faint real signal.

What I'm looking for:

Clean, aligned linear masters without obvious residual gradients or correctable optical blur, while retaining faint low-SNR structure. Avoid preprocessing that makes the data cosmetically smooth at the expense of SII or OIII.

4. EXTRACT H α , OIII, AND SII

Use Seti Astro DBExtract, or another validated dual-band extraction method, on the linear registered masters.

For this workflow:

L-eXtreme master → H α + OIII source data

L-Synergy master → SII + OIII source data

When the extraction tool provides sensor-specific choices, select the configuration that matches the OSC camera used for the project. For the IC 1396 worked example, that was the ASI533/IMX533 configuration.

Generate and preserve separate grayscale emission-line masters. For the IC 1396 worked example:

Elephant_Ha

Elephant_OIII

Elephant_SII

OIII is present in both dual-band datasets. A suitable extraction workflow can make use of information from both rather than automatically discarding one OIII dataset.

What I'm looking for:

Three usable grayscale masters containing distinct, believable nebular structure. At this point, preservation of the extracted line data is more important than how any individual channel looks aesthetically.

5. INSPECT THE EMISSION-LINE MASTERS

Use STF for inspection while the H α , OIII, and SII masters remain linear. Do not permanently stretch them merely to make the channels look equally impressive.

H α :

Expect the strongest signal in many emission nebulae. Inspect the main nebular structure, fine filaments, and—in IC 1396—the trunk and surrounding emission.

OIII:

Look for morphology that differs from H α , particularly around ionization fronts and structures that may later create cyan/blue separation in a synthetic palette.

SII:

Inspect carefully for genuine structure and noise. SII may be the limiting channel. Its quality can determine whether additional L-Synergy acquisition would materially improve the project.

What I'm looking for:

Real structural differences among the three channels, adequate SII signal-to-noise, and no extraction artifact that should be corrected before palette construction.

Decision point:

If SII is too weak or noisy to contribute meaningful structure, consider acquiring more L-Synergy data rather than forcing the channel during later processing. If all three masters contain useful information, proceed to palette construction.

6. CREATE A BASELINE SHO REFERENCE

Before exploring alternative palettes, create and preserve a conventional SHO reference from the extracted emission-line masters:

R = SII

G = H α

B = OIII

This reference is useful because it shows the direct channel assignment and provides a consistent point of comparison. When H α dominates, the initial SHO image may appear strongly green. That is normal. It is a mapped representation of the three channels, not a finished aesthetic.

Keep this reference linear and save it without overwriting the H α , OIII, or SII source masters. Do not automatically apply the normal broadband SPCC workflow to the synthetic SHO combination; the displayed colors are deliberately assigned rather than natural broadband color.

What I'm looking for:

A technically sound reference image in which all three emission-line masters are contributing and no source channel has been lost or clipped. I am not judging the final palette at this stage.

7. DEVELOP THE PALETTE — PERFECT PALETTE PICKER

Treat palette development as its own processing stage. The goal is not to repair the green baseline SHO image. The goal is to use the preserved H α , OIII, and SII information to explore different visual mappings of real spatial differences in the nebula.

Use Seti Astro Perfect Palette Picker with the extracted emission-line masters and evaluate multiple palette interpretations while preserving the source masters and baseline SHO reference. Compare how each treatment separates H α -dominated regions from OIII- and SII-rich structures, especially where the channels show genuinely different morphology.

Foraxx is one possible palette interpretation, not the definition of this workflow. In the IC 1396 project, the Foraxx result became the preferred finished version after multiple palette treatments were evaluated. A different target may favor a different interpretation.

What I'm looking for:

Three to five palette combinations that look most interesting and still preserve useful structural separation among H α , OIII, and SII. Render promising combinations into separate images so they can be viewed at useful scale and carried forward for additional processing before choosing a final interpretation.

DECISION POINT — CHOOSE THE PALETTE THAT SERVES THE DATA

Compare candidate palettes at both fit-to-screen and higher magnification. Ask whether the palette makes meaningful differences in the emission data easier to see or merely produces a more dramatic color scheme.

Pay particular attention to:

- whether weak OIII remains visible rather than being overwhelmed by H α
- whether SII contributes recognizable structure rather than simply shifting the global color
- whether color transitions follow actual nebular morphology
- whether the treatment creates excessive gold/orange saturation or electric cyan/blue
- whether faint structures survive the palette transformation

Render and preserve several promising alternatives. The same H α , OIII, and SII masters can support conventional SHO, Foraxx, and other synthetic narrowband interpretations without restacking the original light frames. Comparing several rendered candidates is often more useful than trying to choose a palette from a small preview alone.

REPRODUCIBILITY NOTE

The exact Perfect Palette Picker / Foraxx settings used during the original Elephant Trunk processing session were not preserved, so this guide does not reconstruct them from memory or present invented values. The reproducible part of the workflow is the architecture: separate complementary-filter integrations, registered linear masters, preserved H α /OIII/SII extractions, and nondestructive palette exploration.

8. VERIFY PPP OUTPUT STATE & CHOOSE A FINISHING APPROACH

At this point, the core complementary dual-band / Perfect Palette Picker workflow is complete. A preferred synthetic narrowband palette has been created from the preserved H α , OIII, and SII masters.

Do not infer the image state from PPP's Linear Input Data setting alone. Immediately after rendering the selected palette, reset or turn off STF and inspect the image with no screen stretch. If the image remains normally visible, treat the PPP output as nonlinear and continue with nonlinear refinement. If it collapses to black or nearly black, treat it as linear and apply a permanent stretch with HistogramTransformation and/or GHS before nonlinear processing. This verification determines the next processing step and avoids either applying nonlinear tools too early or unnecessarily stretching an image that is already nonlinear.

For greater control, my preferred approach is to separate the stars from the nebula and process them independently. The remaining steps document that optional advanced finishing path.

SIMPLIFIED FINISHING PATH

Selected PPP palette → turn STF off/reset it → verify linear/nonlinear state → if linear, apply a permanent stretch with HistogramTransformation and/or GHS; if nonlinear, do not re-stretch → modest color/contrast adjustment → appropriate noise/sharpening work → final quality check.

ADVANCED FINISHING PATH — ASTRO DIVER APPROACH

Selected PPP palette → turn STF off/reset it → verify linear/nonlinear state → StarXTerminator → process starless/stars according to their actual image state → recombine → final quality check.

Decision point:

Choose the finishing path that matches the image, available tools, and desired level of control. Star separation is not required to benefit from the complementary dual-band / PPP workflow. If you want a workflow with more direct control over channel blending and nonlinear development rather than PPP palette selection, see the related Dynamic HOO Processing guide.

9. OPTIONAL ADVANCED FINISHING — STAR SEPARATION

For the advanced path, use StarXTerminator to create a starless nebula image and a corresponding star image, and preserve both. Carry forward the image-state decision made immediately after PPP: if the selected palette was linear, it still requires a permanent stretch; if it was already nonlinear, do not apply another initial stretch.

Use palette-neutral file names unless the image is intentionally being preserved as a specific palette version. For example:

SelectedPalette_starless

SelectedPalette_stars

Processing the two components independently allows stronger color and contrast development in the nebula without forcing the stars through the same transformations.

What I'm looking for:

Clean removal of stars without obvious holes, ringing, halos, or damaged fine nebular structure. Inspect bright stars and dense star fields closely before continuing.

10. STRETCH AND DEVELOP THE STARLESS NEBULA

Develop the starless image according to the verified PPP output state. If it is still linear, Generalized Hyperbolic Stretch is especially useful for a controlled first permanent stretch while protecting bright structures and bringing faint emission forward. If it is already nonlinear, skip the initial stretch and move directly to restrained tonal, color, and local-contrast refinement.

When a permanent stretch is required, use restrained passes rather than attempting the entire stretch at once. Then use CurvesTransformation, ColorSaturation, masks, and other targeted tools as needed to refine separation among structures represented by H α , OIII, and SII.

For IC 1396, the goals included revealing faint outer emission, preserving the dark Elephant Trunk, retaining structure in brighter nebulosity, and avoiding a crushed background.

Do not treat removal of green—or production of any particular color—as the objective. The selected PPP palette has already established the basic color interpretation. Finishing should improve visual separation, local contrast, and balance while preserving the underlying emission morphology.

What I'm looking for:

Faint structure becoming visible without clipped blacks or blown highlights, coherent color relationships, and preservation of weak OIII and SII rather than allowing stronger H α to overwhelm them.

11. PROCESS THE STARS

Process the star image separately and comparatively conservatively. Stretch the stars with an appropriate HistogramTransformation or GHS approach and adjust saturation only as needed.

Stars derived from narrowband data may not have natural broadband color. Avoid forcing them to look like broadband stars through aggressive color manipulation. If natural stellar color is important, a separately acquired broadband or UV/IR-cut star dataset can be substituted when available.

What I'm looking for:

Stars that support the nebula rather than dominate it—controlled size, restrained saturation, no unnecessary clipping, and no obvious processing artifacts.

12. RECOMBINE STARS AND STARLESS IMAGE

Recombine the processed stars with the finished starless nebula using ImageBlend or another controlled recombination method appropriate to the data.

Judge the recombined image at both fit-to-screen and higher magnification. If the stars suddenly dominate the composition, appear too saturated, or expose halos and StarX artifacts that were less obvious separately, return to the star layer rather than compensating by damaging the nebula.

What I'm looking for:

A recombination that feels integrated. The nebula should retain the palette and contrast developed in the starless stage while the stars return naturally to the field.

13. FINAL QUALITY CHECK

Whether using the simplified or advanced finishing path, inspect the full frame and several 100% regions before calling the image finished.

Check for:

- clipped or unnaturally empty background
- excessive gold/orange saturation
- electric, disconnected, or unnaturally dominant OIII color
- residual green that distracts from the chosen palette rather than representing intentional structure
- magenta, oversized, or excessively saturated stars
- halos, holes, ringing, or other star-removal artifacts if star separation was used
- oversharpened filaments or artificial local-contrast edges
- excessive noise reduction, particularly where SII or OIII is weak
- residual gradients or uneven background
- color transitions that do not correspond to believable emission structure

The final question is not whether the image matches a predetermined palette appearance. It is whether the selected interpretation preserves the useful information in the H α , OIII, and SII data and presents it coherently.

14. FILES TO PRESERVE

Keep the independent filter masters, extracted emission-line masters, baseline SHO combination, preferred palette combination, starless/stars outputs and final processed image. The H α , OIII, and SII masters are especially valuable because alternate palettes can be developed later without restacking the original lights.

WORKED EXAMPLE — IC 1396 / ELEPHANT TRUNK

IC 1396 was the learning project for this complementary dual-band OSC approach.

Equipment and documented starting data:

Telescope: Askar 111F APO

Camera: ZWO ASI533MC Pro cooled one-shot-color camera

Mount: ZWO AM5

L-eXtreme: 244 × 90 seconds = 6 hours 6 minutes

L-Synergy: complementary SII + OIII dataset acquired separately

The original project plan was straightforward: integrate the two filters independently, align the masters, extract H α /OIII/SII, create a conventional SII → R, H α → G, OIII → B SHO image, and then develop the color manually.

The actual processing session changed the emphasis of the workflow. Seti Astro DBExtract became the practical route to the emission-line masters, and Perfect Palette Picker made it possible to evaluate multiple synthetic narrowband interpretations from those preserved channels. The Foraxx interpretation ultimately became the preferred finished IC 1396 image.

That change produced the most important lesson from the project: the acquisition architecture is more valuable than committing to a palette in advance. Separate complementary-filter integrations and preserved H α , OIII, and SII masters allow the same data to be revisited as

conventional SHO, Foraxx, or other synthetic narrowband interpretations without recalibrating or restacking the original light frames.

As noted above, the exact PPP / Foraxx settings from the original session were not preserved. The worked example therefore focuses on the repeatable architecture and decision process.

LESSONS LEARNED

- Integrate complementary filters separately.
- Keep the data linear through line extraction and initial palette construction.
- Preserve H α , OIII, and SII masters permanently.
- Protect weak SII during denoising and stretching.
- Treat palette development as interpretation of the line data, not merely removal of green.
- Star separation is optional. Use it when independent control of the nebula and stars materially improves the finishing process.
- Use settings as image-specific decisions rather than universal recipes.
- Record exact PPP, GHS, Curves, and recombination settings during future runs when they materially affect reproducibility.

ABOUT THIS GUIDE

This guide documents the workflow and decisions used in the Astro Diver IC 1396 project while separating repeatable technique from image-specific choices. Exact settings that were not preserved during the original processing session are intentionally omitted rather than reconstructed from memory. Use the sequence and decision points as a framework, then adjust settings to the quality and characteristics of your own data.